

MAGIC TeV Gamma-Ray Observations of Markarian 421 during Multiwavelength Campaigns in 2006

J. Aleksić¹, H. Anderhub², L. A. Antonelli³, P. Antoranz⁴, M. Backes⁵, C. Baixeras⁶, S. Balestra⁴, J. A. Barrio⁴, D. Bastieri⁷, J. Becerra González⁸, J. K. Becker⁵, W. Bednarek⁹, A. Berdyugin¹⁰, K. Berger⁹, E. Bernardini¹¹, A. Biland², R. K. Bock^{12,7}, G. Bonnoli¹³, P. Bordas¹⁴, D. Borla Tridon¹², V. Bosch-Ramon¹⁴, D. Bose⁴, I. Braun², T. Bretz¹⁵, D. Britzger¹², M. Camara⁴, E. Carmona¹², A. Carosi³, P. Colin¹², S. Commichau², J. L. Contreras⁴, J. Cortina¹, M. T. Costado^{8,16}, S. Covino³, F. Dazzi^{17,26}, A. De Angelis¹⁷, E. de Cea del Pozo¹⁸, R. De los Reyes^{4,28}, B. De Lotto¹⁷, M. De Maria¹⁷, F. De Sabata¹⁷, C. Delgado Mendez^{8,27}, M. Doert⁵, A. Domínguez¹⁹, D. Dominis Prester²⁰, D. Dorner², M. Doro⁷, D. Elsaesser¹⁵, M. Errando¹, D. Ferenc²¹, M. V. Fonseca⁴, L. Font⁶, R. J. García López^{8,16}, M. Garczarczyk⁸, M. Gaug⁸, N. Godinovic²⁰, D. Hadasch¹⁸, A. Herrero^{8,16}, D. Hildebrand², J. Hose¹², D. Hrupec²⁰, C. C. Hsu¹², T. Jogler¹², S. Klepser¹, T. Krähenbühl², D. Kranich², A. La Barbera³, A. Laille²¹, E. Leonardo¹³, E. Lindfors¹⁰, S. Lombardi⁷, F. Longo¹⁷, M. López⁷, E. Lorenz^{2,12}, P. Majumdar¹¹, G. Maneva²², N. Mankuzhiyil¹⁷, K. Mannheim¹⁵, L. Maraschi³, M. Mariotti⁷, M. Martínez¹, D. Mazin¹, M. Meucci¹³, J. M. Miranda⁴, R. Mirzoyan¹², H. Miyamoto¹², J. Moldón¹⁴, M. Moles¹⁹, A. Moralejo¹, D. Nieto⁴, K. Nilsson¹⁰, J. Ninkovic¹², R. Orito¹², I. Oya⁴, R. Paoletti¹³, J. M. Paredes¹⁴, S. Partini¹³, M. Pasanen¹⁰, D. Pascoli⁷, F. Pauss², R. G. Pegna¹³, M. A. Perez-Torres¹⁹, M. Persic^{17,23}, L. Peruzzo⁷, F. Prada¹⁹, E. Prandini⁷, N. Puchades¹, I. Puljak²⁰, I. Reichardt¹, W. Rhode⁵, M. Ribó¹⁴, J. Rico^{24,1}, M. Rissi², S. Rügamer¹⁵, A. Saggion⁷, T. Y. Saito¹², M. Salvati³, M. Sánchez-Conde¹⁹, K. Satalecka¹¹, V. Scalzotto⁷, V. Scapin¹⁷, T. Schweizer¹², M. Shayduk¹², S. N. Shore²⁵, A. Sierpowska-Bartosik⁹, A. Sillanpää¹⁰, J. Sitarek^{12,9}, D. Sobczynska⁹, F. Spanier¹⁵, S. Spiro³, A. Stamerra¹³, B. Steinke¹², N. Strah⁵, J. C. Struebig¹⁵, T. Suric²⁰, L. Takalo¹⁰, F. Tavecchio³, P. Temnikov²², D. Tescaro¹, M. Teshima¹², D. F. Torres^{24,18}, H. Vankov²², R. M. Wagner¹², V. Zabalza¹⁴, F. Zandanel¹⁹, and R. Zanin¹

(Affiliations can be found after the references)

Received 2009 December 23

ABSTRACT

Context. Wide-range spectral coverage of blazar-type active galactic nuclei is of paramount importance for understanding the particle acceleration mechanisms assumed to take place in their jets. The Major Atmospheric Gamma Imaging Cerenkov (MAGIC) telescope participated in three multiwavelength (MWL) campaigns, observing the blazar Markarian (Mkn) 421 during the nights of 2006 April 28, 29, and 2006 June 14.

Aims. We analyzed the corresponding MAGIC very-high energy observations during 9 nights from 2006 April 22 to 30 and on 2006 June 14. We inferred light curves with sub-day resolution and night-by-night energy spectra.

Methods. MAGIC detects γ -rays by observing extended air showers in the atmosphere. The obtained air-shower images were analyzed using the standard MAGIC analysis chain.

Results. A strong γ -ray signal was detected from Mkn 421 on all observation nights. The flux ($E > 250$ GeV) varied on night-by-night basis between $(0.92 \pm 0.11) 10^{-10} \text{ cm}^{-2} \text{ s}^{-1}$ (0.57 Crab units) and $(3.21 \pm 0.15) 10^{-10} \text{ cm}^{-2} \text{ s}^{-1}$ (2.0 Crab units) in 2006 April. There is a clear indication for intra-night variability with a doubling time of $36 \pm 10_{\text{stat}}$ minutes on the night of 2006 April 29, establishing once more rapid flux variability for this object. For all individual nights γ -ray spectra could be inferred, with power-law indices ranging from 1.66 to 2.47. We did not find statistically significant correlations between the spectral index and the flux state for individual nights. During the June 2006 campaign, a flux substantially lower than the one measured by the Whipple 10-m telescope four days later was found. Using a log-parabolic power law fit we deduced for some data sets the location of the spectral peak in the very-high energy regime. Our results confirm the indications of rising peak energy with increasing flux, as expected in leptonic acceleration models.

Key words. Gamma rays: galaxies – BL Lacertae objects: individual (Mkn 421) – Radiation mechanisms: non-thermal

1. Introduction

The active galactic nucleus (AGN) Markarian (Mkn) 421 was the first extragalactic source detected in the TeV energy range, using imaging atmospheric Cerenkov telescopes (IACTs; Punch et al., 1992; Petry et al., 1996). With a redshift of $z = 0.030$ it is the closest known and, along with Mkn 501, the best-studied

TeV γ -ray emitting blazar.¹ So far, flux variations by more than one order of magnitude (e.g., Fossati et al., 2008), and occasional flux doubling times as short as 15 min (Gaidos et al., 1996; Aharonian et al., 2002; Schweizer, Wagner & Lorenz, 2008) have been observed. Variations in the hardness of the TeV γ -ray spectrum during flares were reported by sev-

Send offprint requests to: snruegam@astro.uni-wuerzburg.de, oya@gae.ucm.es, robert.wagner@mpp.mpg.de

¹ See, e.g., <http://www.mpp.mpg.de/~rwagner/sources/> for an up-to-date list of VHE γ -ray sources.

eral groups (e.g. Krennrich et al., 2002; Aharonian et al., 2005; Fossati et al., 2008). Simultaneous observations in the X-ray and very-high energy (VHE; $E \gtrsim 100$ GeV) bands show strong evidence for correlated flux variability (Krawczynski et al., 2001; Błażejowski et al., 2005; Fossati et al., 2008). With a long history of observations, Mkn 421 is an ideal candidate for long-term and statistical studies of its emission (Tluczykont et al., 2007; Goebel et al., 2008a; Hsu et al., 2009).

Mkn 421 has been detected and studied at basically all wavelengths of the electromagnetic spectrum from radio waves up to VHE γ -rays. Its wide-range spectral energy distribution (SED) shows the typical double-peak structure of AGN. Mkn 421 is a so-called blazar. These constitute a rare subclass of AGNs with beamed emission closely aligned to our line of sight. In blazars, the low-energy peak at keV energies is thought to arise dominantly from synchrotron emission of electrons, while the origin of the high-energy (GeV-TeV) bump is still debated. The SED is commonly interpreted as being due to the beamed, non-thermal emission of synchrotron and inverse-Compton radiation from ultrarelativistic electrons. These are assumed to be accelerated by shocks moving along the jets at relativistic bulk speed. For most of the observations, the SED can be reasonably well described by homogeneous one-zone synchrotron-self-Compton (SSC) models (e.g. Marscher & Gear, 1985; Maraschi et al., 1992; Costamante & Ghisellini, 2002). Hadronic models (Mannheim et al., 1996; Mücke et al., 2003), however, can also explain the observed features. A way to distinguish between the different emission models is to determine the positions, evolution and possible correlations (see, e.g., Wagner, 2008b, for a review) of both peaks in the SED, using simultaneous, time-resolved data covering a broad energy range, e.g., as obtained in multiwavelength (MWL) observational campaigns.

In this Paper we present results from Major Atmospheric Gamma-ray Imaging Cerenkov (MAGIC) telescope VHE γ -ray observations of Mkn 421 during eight nights from 2006 April 22 to 30, and on 2006 June 14. For most of the days, optical R -band observations were conducted with the KVA telescope. Simultaneous observations were carried out by *Suzaku* (Mitsuda et al., 2007) and H.E.S.S., as well as by *XMM-Newton* (Jansen et al., 2001) on 2006 April 28 and 29, respectively. During both nights, we carried out particularly long, uninterrupted observations in the VHE energy band of ≈ 3 h duration each. An onset of activity in the X-ray band triggered an *INTEGRAL*-led target-of-opportunity (ToO) campaign, which took place from 2006 June 14 – 25 for a total of 829 ks (Lichti et al., 2008). Within this campaign, MAGIC observed Mkn 421 at rather high zenith angles from 43 to 52 degrees in parallel with *INTEGRAL* on 2006 June 14.

In the following sections, we describe the data sets and the analysis applied to the VHE γ -ray data, the determination of spectra for all observation nights, and put the results into perspective with other VHE γ -ray observations of Mkn 421. The interpretation of these data in a MWL context is carried out in Acciari et al. (2009) and subsequent papers.

VHE γ -ray observations in 2006 April and June have also been carried out by the Whipple telescope (Horan et al., 2009), by the VERITAS (Fegan, 2008), and TACTIC (Yadav et al., 2007) collaborations, although not simultaneously with our observations.

2. The MAGIC telescope

The VHE γ -ray observations were conducted with the MAGIC telescope located on the Canary island La Palma (2200 m above

Table 1. Some characteristic parameters of the different data sets of the campaign.

Night	Observation Window [MJD]	$t_{\text{eff.}}$ [h]	ZA [°]
2006/04/22	53847.97679 – 53848.01460	0.76	18 – 28
2006/04/24	53849.96428 – 53850.00669	0.99	16 – 28
2006/04/25	53850.92813 – 53850.99607	1.53	10 – 26
2006/04/26	53851.92862 – 53852.00383	1.64	10 – 29
2006/04/27	53852.93474 – 53853.00047	1.42	12 – 28
2006/04/28	53853.88173 – 53854.01394	2.23	10 – 32
2006/04/29	53854.89514 – 53855.04119	2.78	9 – 41
2006/04/30	53855.97283 – 53855.97906	0.16	23 – 24
2006/06/14	53900.91979 – 53900.95532	0.80	43 – 52

$t_{\text{eff.}}$ denotes the effective observation time. ZA gives the zenith angle range of the observations.

sea level, 28°45'N, 17°54'W). At the time of our observations in 2006, MAGIC was a single-dish 17-m \emptyset instrument² for the detection of atmospheric air showers induced by γ -rays. Its hexagonally-shaped camera with a field of view (FOV) of $\approx 3.5^\circ$ mean diameter comprises 576 high-sensitivity photomultiplier tubes (PMTs): 180 pixels of $0.2^\circ \emptyset$ surround the inner section of the camera of 394 pixels of $0.1^\circ \emptyset$ ($= 2.2^\circ \emptyset$ FOV). The trigger is formed by a coincidence of ≥ 4 neighboring pixels. Presently the accessible trigger energy range (using the MAGIC standard trigger; Meucci et al. 2007) spans from 50 – 60 GeV (at small zenith angles) up to tens of TeV. Further details, telescope parameters, and performance information can be found in Baixeras et al. (2004); Cortina et al. (2005); Albert et al. (2008a).

3. Observations and data analysis

The observations were carried out during dark nights, employing the so-called wobble mode (Daum et al., 1997), in which two opposite sky directions, each 0.4° off the source, are tracked alternately for 20 minutes each. The on-source data are defined by calculating image parameters with respect to the source position, whereas background control (“off”) data are obtained from the same data set, but with image parameters calculated with respect to the position on the opposite side of the camera, the *anti-source* position. The simultaneous measurement of signal and background makes additional background control data unnecessary. In order to avoid an unwanted contribution from γ -events in the off sample, and to guarantee the statistical independence between the on and the off samples in the signal region, events included in the signal region of the on sample were excluded from the off sample and vice versa.

The data were analyzed following the standard MAGIC analysis procedure (Bretz & Wagner, 2003; Bretz & Dorner, 2008). After calibration (Albert et al., 2008c) and extracting the signal at the pulse maximum using a spline method, the air-shower images were cleaned of noise from night-sky background light by applying a three-stage image cleaning. The first stage requires a minimum number of 6 photoelectrons in the core pixels and 3 photoelectrons in the boundary pixels of the images (see, e.g. Fegan, 1997). These tail cuts are scaled according to the larger size of the outer pixels of the MAGIC camera. Only pixels with at least two adjacent pixels with a signal arrival time difference lower than 1.75 ns survive the second cleaning stage. The third stage repeats the cleaning of the second stage, but requires only one adjacent pixel within the 1.75 ns time window.

² Since 2009, MAGIC is a two-telescope stereoscopic system (Cortina, Goebel & Schweizer, 2009).

The data were filtered by rejecting trivial background events, such as accidental noise triggers, triggers from nearby muons, or data taken during adverse atmospheric conditions (e.g., low atmospheric transmission). 12.7 hours out of the total 15.0 hours' worth of data survived the latter quality selection and were used for further analysis.

We calculated image parameters (Hillas, 1985) such as WIDTH, LENGTH, SIZE, CONC, M3LONG (the third moment of the light distribution along the major image axis), and LEAKAGE (the fraction of light contained in the outermost ring of camera pixels) for the surviving events. For the γ /hadron separation, a SIZE-dependent parabolic cut in AREA \equiv WIDTH \times LENGTH $\times \pi$ was used (Riegel et al., 2005). The cut parameters for the assessment of the detection significance were optimized on Mkn 421 data from close-by days. For the data of the 2006 June 14 at rather large zenith angles, data of Mkn 501 from 2006 October were used to determine the optimal cuts. Any significance in this work was calculated using Eq. 17 of Li & Ma (1983).

The primary γ -ray energies were reconstructed from the image parameters using a Random Forest regression method (Albert et al., 2008b, and references therein) trained with Monte-Carlo simulated events (MCs; Knapp & Heck, 2004; Majumdar et al., 2005). The MC sample is characterized by a power-law spectrum between 10 GeV and 30 TeV with a differential spectral photon index of $\alpha = -2.6$, and a point-spread function resembling the experimental one. The events were selected to cover the same zenith distance range as the data. For the spectrum calculation, the area cut parameters were optimized to yield a constant MC cut efficiency of 90% over the whole energy range, increasing the γ -ray event statistics at the threshold.

The Mkn 421 observations presented here are among the first data taken by MAGIC after major hardware updates in April 2006 (Goebel et al., 2008b), which required us to thoroughly examine the data. Despite the hardware changes, the MAGIC subsystems performed as expected with the exception of an unstable trigger behavior for some PMTs, leading to a significant loss of events in one of the six sectors of the camera. In order to proceed with the data analysis with serenity and to estimate the effect caused by this inhomogeneity, a simple procedure was applied to the data: The expected number of events, as a function of energy, for the affected sector was estimated as the mean of the number of events in the other five sectors of the camera. (A homogeneous distribution of events through the six sectors is expected for normal conditions). The difference was computed using the whole data sample in order to have sufficient statistics. We found a decrease of the differential photon flux of 5.7% between 250 and 400 GeV, 4.6% between 400 and 650 GeV, 2.2% between 650 and 1050 GeV and $< 1\%$ for higher energies for the 2006 April data. Due to the higher zenith distance and energy threshold, the method was adapted for 2006 June 14 and yielded a decrease of 5.2% between 450 and 670 GeV and 2.6% for higher energies. However the above mentioned effect is just an average one, with estimated flux errors of up to 6.6% showing up for individual nights.

To mitigate the effect of the inhomogeneity, instead of an (already increased) energy threshold of 250 GeV, higher thresholds of 350 or 450 GeV were applied for some observation nights. In this way we made sure that the estimated systematic error remains within reasonable limits.

For the calculation of the individual light curves as well as for the overall 2006 April lightcurve, the flux between 250 GeV and 350 GeV was extrapolated for the nights with higher threshold. We assumed a power-law behavior in this energy range, with

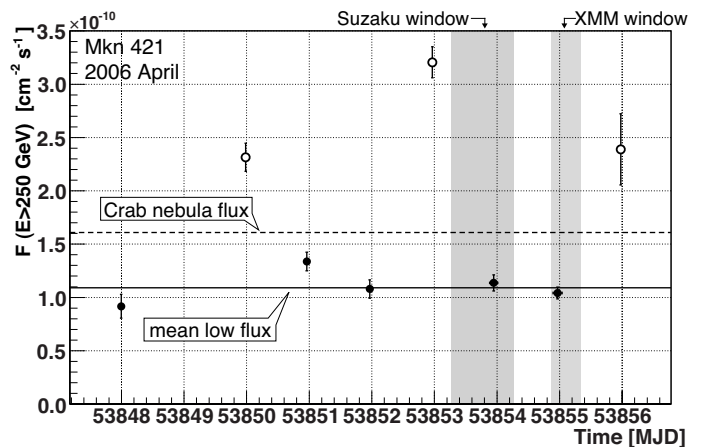


Fig. 1. VHE ($E > 250$ GeV) light curve for Mkn 421 observations in April 2006. The data points represent average nightly fluxes. The observation windows of the *Suzaku* (MJD 53853.28–53854.27) and *XMM-Newton* (MJD 53854.87–53855.35) MWL campaigns are marked by the gray-shaded areas. A “mean low flux” (solid line) was averaged over all data points below $1.6 \cdot 10^{-10} \text{ cm}^{-2} \text{ s}^{-1}$, i.e., excluding those data points marked by thin open circles. The dashed line gives the Crab nebula flux (Albert et al., 2008a) for comparison.

the spectral index determined for the first three energy bins of the whole April dataset (i.e., $\alpha = -2.08$). The flux normalization for each night has been determined at 500 GeV by a fit to the first three differential spectral points, an energy range which is reliable for all affected nights.

Tab. 1 summarizes the analyzed data sets. The statistical significance of any detection is assessed by applying a cut in θ^2 , where θ is the angular distance between the expected source position and the reconstructed γ -ray arrival direction. The arrival directions of the showers in equatorial coordinates were calculated using the DISP method (Fomin et al., 1994; Lessard et al., 2001). We replaced the constant coefficient ξ in the parameterization of DISP in the original approach by a term which is dependent on LEAKAGE, SIZE, and SLOPE,

$$\xi = \xi_0 + \xi_1 \text{SLOPE} + \xi_2 \text{LEAKAGE} + k\xi_3 (\log_{10} \text{SIZE} - \xi_4)^2, \quad (1)$$

$k = 0$ for $\log_{10} \text{SIZE} < \xi_4$ and $k = 1$ for $\log_{10} \text{SIZE} \geq \xi_4$. The coefficients were determined using simulated data. The parameter SLOPE is a measure for the longitudinal arrival time evolution of the shower in the camera plane similar to the time parameter GRADIENT in Aliu et al. (2009). Instead of defining the parameter from a fit to the arrival time distribution, however, SLOPE is determined as an analytical solution of the fit. Note that this new parametrization makes DISP and therefore θ^2 source dependent.

All stated errors are statistical errors only; we estimate our systematic errors to be 16% for the energy scale, 11% for absolute fluxes and flux normalizations, and 0.2 for the spectral slopes (Albert et al., 2008a), not including the additional systematic flux errors mentioned above.

A second, independent analysis of the data yielded compatible results to those presented here.

4. Results

4.1. Results for 2006 April 22 – 30

MAGIC observed Mkn 421 from MJD 53847 to MJD 53855. During the observations, two MWL campaigns were carried out

Table 2. Analysis results.

Observation Night	N_{excess}	$N_{\text{backgr.}}$	S	E_{cut} [GeV]	$F(E > E_{\text{min}})$	$\chi^2_{\text{red,const}}$	f_0	α	$\chi^2_{\text{red,PL}}$
2006/04/22	100	29	10.9σ	350	0.92 ± 0.11	1.3/2	0.98 ± 0.15	2.05 ± 0.21	2.1/2
2006/04/24	419	69	25.0σ	250	2.32 ± 0.13	2.7/2	2.45 ± 0.14	2.25 ± 0.09	2.0/3
2006/04/25	342	83	20.8σ	250	1.34 ± 0.09	1.7/2	1.43 ± 0.09	2.26 ± 0.12	0.24/3
2006/04/26	225	62	16.4σ	350	1.08 ± 0.09	1.3/4	1.21 ± 0.11	2.35 ± 0.17	0.41/2
2006/04/27	615	56	33.5σ	350	3.21 ± 0.15	1.9/4	3.37 ± 0.18	2.07 ± 0.07	4.8/4
2006/04/28	311	75	19.9σ	350	1.14 ± 0.08	4.3/8	1.32 ± 0.10	2.47 ± 0.14	0.65/2
2006/04/29	514	169	23.7σ	250	1.04 ± 0.06	41/7	1.14 ± 0.06	2.28 ± 0.09	2.0/4
2006/04/30	69	11	10.3σ	250	2.39 ± 0.33	—	2.16 ± 0.34	1.66 ± 0.20	1.4/1
2006/06/14	95	87	7.5σ	450	0.34 ± 0.06	2.4/1	0.168 ± 0.032	2.38 ± 0.44	1.5/2

Number of excess (N_{excess}) and background ($N_{\text{backgr.}}$) events, resulting significances S , lower cuts in event energy, integral fluxes F above $E_{\text{min}} = 250$ GeV for the 2006 April data and $E_{\text{min}} = 450$ GeV for the 2006 June 14 data (in units of $10^{-10} \text{ cm}^{-2} \text{ s}^{-1}$), fit quality of a constant-flux fit to the individual observation nights (see Fig. 2), and power-law fit results for the differential energy spectra of $dF/dE = f_0 \cdot (E/E_0)^{-\alpha}$ with $E_0 = 0.5$ TeV for the 2006 April data and $E_0 = 1.0$ TeV for the 2006 June 14 data, respectively; f_0 in units of $10^{-10} \text{ TeV}^{-1} \text{ cm}^{-2} \text{ s}^{-1}$.

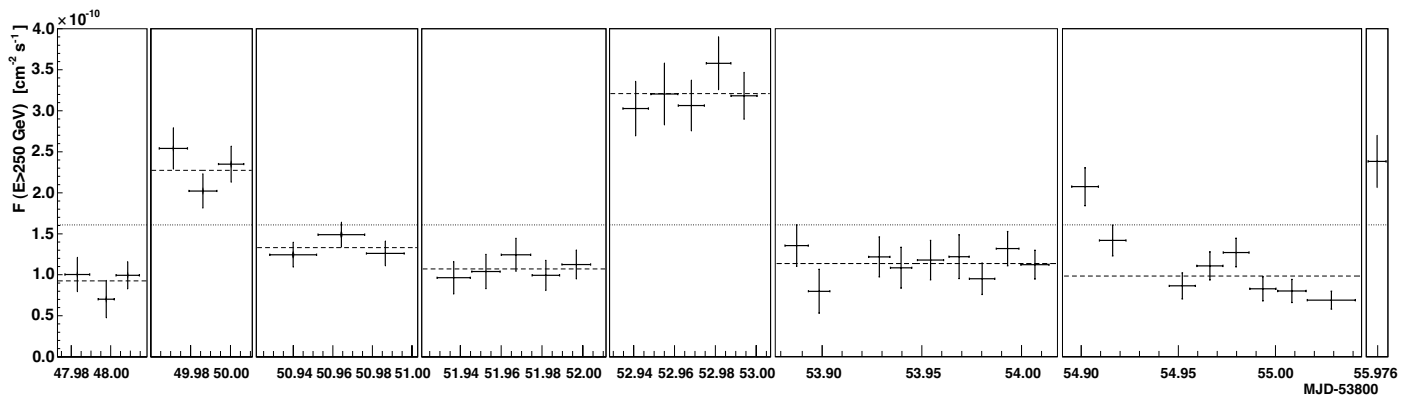


Fig. 2. VHE ($E > 250$ GeV) light curve for Mkn 421 observations in April 2006. The dotted line represents the Crab nebula flux (Albert et al., 2008a), whereas the individual dashed lines show the result of a fit to the time bins (average nightly flux) of the corresponding nights.

simultaneously with *Suzaku* and with *XMM-Newton* on MJD 53854 and MJD 53855, respectively. Mkn 421 was also observed as part of the monitoring program of the Whipple 10-m telescope (see Horan et al., 2009), albeit about 3.5 hours after the MAGIC observations stopped, due to the different longitudes of the two instruments.

A strong γ -ray signal from the source was detected in all eight observation nights. In total, 3165 excess events were recorded over a background of 693 events, yielding an overall significance of 64.8σ . Mkn 421 exhibited an average flux of $F_{>250 \text{ GeV}} = (1.48 \pm 0.03) \cdot 10^{-10} \text{ cm}^{-2} \text{ s}^{-1}$. When compared to earlier observations (see, e.g. Albert et al., 2007a; Gluzykont et al., 2007; Goebel et al., 2008a; Steele et al., 2008), our observations indicate an elevated flux state of Mkn 421. We found high flux states in the nights of MJD 53850, $F_{>250 \text{ GeV}} = (2.32 \pm 0.13) \cdot 10^{-10} \text{ cm}^{-2} \text{ s}^{-1}$, MJD 53853, $F_{>250 \text{ GeV}} = (3.21 \pm 0.15) \cdot 10^{-10} \text{ cm}^{-2} \text{ s}^{-1}$, and MJD 53856, $F_{>250 \text{ GeV}} = (2.39 \pm 0.33) \cdot 10^{-10} \text{ cm}^{-2} \text{ s}^{-1}$ (Fig. 1). In the remaining nights (we assumed nights with fluxes below $1.6 \cdot 10^{-10} \text{ cm}^{-2} \text{ s}^{-1}$ as non-flare nights), Mkn 421 exhibited a low-flux average of $F_{>250 \text{ GeV}} = (1.09 \pm 0.03) \cdot 10^{-10} \text{ cm}^{-2} \text{ s}^{-1}$. The analysis results on a night-by-night basis are summarized in Tab. 2, and include the nightly numbers for excess and background events, significances, and average integral fluxes above 250 GeV (where the nights with an energy cut of 350 GeV where extrapolated down to 250 GeV, see Sect. 3 for

details). The results of a spectral fit based on a simple power law (PL) of the form

$$\frac{dF}{dE} = f_0 \cdot 10^{-10} \text{ TeV}^{-1} \text{ cm}^{-2} \text{ s}^{-1} \left(\frac{E}{E_0} \right)^{-\alpha} \quad (2)$$

are also shown.

The energy thresholds of the individual observations are also given in Tab. 2. As the analysis threshold is always lower than the applied energy cut, the latter one defines the energy threshold value.

The strong γ -ray signal allowed to infer light curves with a resolution below one hour for all of the observation nights, which are shown in Fig. 2 (see Tab. 4³ for the light curve data). Most light curves are compatible with a constant flux during the nightly observation time (see Tab. 2 for all constant-fit χ^2_{red} values), while on MJD 53855 a clear intra-night variability is apparent. A fit with a constant function yields an unacceptable $\chi^2_{\text{red}} = 41/7$ ($P \approx 8 \cdot 10^{-5}\%$) for this night, and the data suggest a flux halving time of $36 \pm 10_{\text{stat}}$ minutes. Note that this interesting observation window has also been covered by *XMM-Newton* observations in the X-ray band (Acciari et al., 2009).

³ Table 4 is available only in the electronic edition of the journal, www.aanda.org

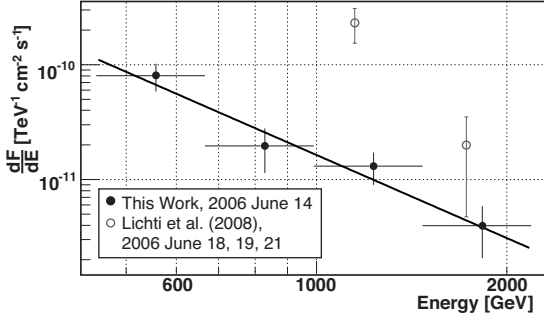


Fig. 3. Differential photon spectrum for Mkn 421 for the observation night of 2006 June 14 (black data points). A power-law fit to the spectrum results in a spectral slope of $\alpha = -2.38 \pm 0.44$ (See Tab. 2 for the fit results). Also shown are spectral points measured with the Whipple 10-m telescope (Lichti et al., 2008) during 2006 June 18-21.

4.2. Results for 2006 June 14

An onset of activity to ≈ 2 times the average quiescent-flux level of Mkn 421 was measured in April 2006 by the *RXTE* all-sky monitor (ASM) instrument. It triggered an *INTEGRAL* ToO campaign from 2006 June 14 to 25 for a total of 829 ks (Lichti et al., 2008). This > 30 mCrab flux remained until September 2006. During the 9-day campaign, Mkn 421 was targeted by various instruments in the radio, optical, X-ray and VHE wavebands. Results are reported in Lichti et al. (2008). On 2006 June 14, MAGIC observed Mkn 421 at rather high zenith angles in parallel with the OMC, JEM-X, and IBIS measurements aboard *INTEGRAL*. Further VHE coverage was provided by the Whipple 10-m telescope on 2006 June 18, 19, and 21 (Lichti et al., 2008).

The MAGIC observations on 2006 June 14 lasted for ≈ 50 minutes. The high zenith angles of 43 to 52 degrees of this observations and the previously mentioned inhomogeneities result in an energy threshold of $E_{\text{thresh.}} = 450$ GeV. In spite of the overall rather difficult observational circumstances caused by the high zenith angle observations (Tonello, 2006; Albert et al., 2006), a firm detection on the $7.5\text{-}\sigma$ significance level was achieved.

The corresponding differential energy spectrum is shown in Fig. 3. Between 450 GeV and 2.2 TeV, it can be described by a simple power-law of the form

$$\frac{dF}{dE} = (1.68 \pm 0.32) \cdot 10^{-11} \text{ TeV}^{-1} \text{ cm}^{-2} \text{ s}^{-1} \left(\frac{E}{1.0 \text{ TeV}} \right)^{-2.38 \pm 0.44} \quad (3)$$

For comparison we also show the spectral points reported by the Whipple 10-m telescope averaged over the nights of 2006 June 18, 19, and 21. Generally, there might be systematic differences between the Whipple and MAGIC measurements. It could, however, be shown that such inter-instrument systematic effects are rather small and under control, e.g. those between MAGIC and H.E.S.S. (Mazin et al., 2005). Particularly the Crab nebula spectra measured by Whipple and MAGIC agree quite well (Albert et al., 2008a). The Mkn 421 flux measured by the Whipple 10-m telescope four days after the MAGIC observation is substantially higher than our measurements (Fig. 3), pointing to a clear evolution of the source emission level within the *INTEGRAL* campaign.

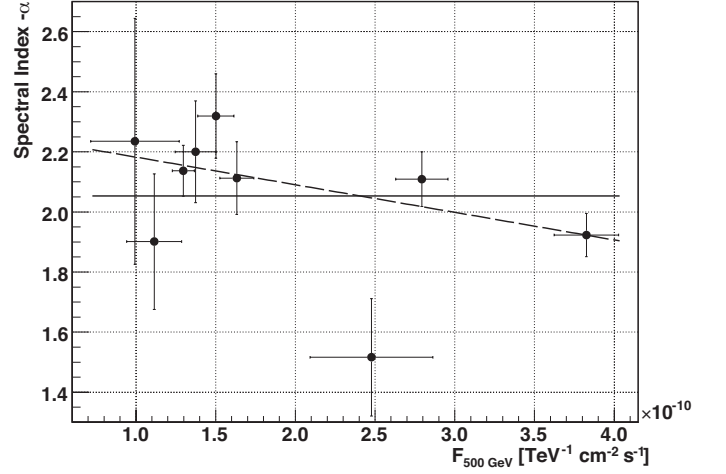


Fig. 4. Spectral index vs. flux at 0.5 TeV deduced from a simple power-law fit after EBL de-absorption for Mkn 421 in 2006 April. The χ^2_{red} for a constant fit (spectral index uncorrelated with flux level; solid line) amounts to $17/8$ ($P \approx 3\%$), while a linear correlation (dashed line) has a $\chi^2_{\text{red}} = 11/7$, equal to $P \approx 12\%$.

5. Discussion

In leptonic acceleration models, e.g., SSC models, a shift of the high-energy peak (attributed to Inverse Compton radiation) in the spectral energy distribution towards higher energies with an increasing flux level is expected. In the VHE domain, such a shift can be traced by spectral hardening. Variations in the hardness of the TeV γ -ray spectrum during flares were reported by several groups (e.g., Krennrich et al., 2002; Aharonian et al., 2005; Fossati et al., 2008). We tested for a correlation of the spectral hardness with the flux level of the de-absorbed spectrum (i.e. after removing any attenuation effects caused by the Extragalactic Background Light [EBL], cf. Nikishov 1962; Gould & Schröder 1966; Hauser & Dwek 2001) in our data (Fig. 4), but found that the correlation neither can be described by a constant fit ($\chi^2_{\text{red}} = 17/8$, $P \approx 3\%$) nor by a linear dependence of spectral hardness and flux level ($\chi^2_{\text{red}} = 11/7$, $P \approx 12\%$), giving no clear preference for either. Although clear flux variations are present in the data set, the overall dynamical range of 3.9 in flux might be too small to see a significant spectral hardening with increasing flux.

The individual night-by-night spectra during the campaign in April 2006 are shown in Fig. 7. All spectral data points are summarized in Tab. 6.⁴ For the nights of 2006 April 22, 26, and 29, there seems to be evidence for a resolved peak, but a likelihood ratio test (e.g., Mazin & Goebel, 2007) yields significant curvature only for 2006 April 27.⁵ We used a logarithmic curvature term, corresponding to a parabolic power-law (log-P) in a $\log(E^2 dF/dE)$ vs. $\log E$ representation (Massaro et al., 2004), and a power-law with exponential cutoff (PL+C) of the form

$$\frac{dF}{dE} = f_0 \cdot 10^{-11} \text{ TeV}^{-1} \text{ cm}^{-2} \text{ s}^{-1} \left(\frac{E}{E_0} \right)^{-(\alpha + \beta \log_{10}(\frac{E}{E_0}))} \quad (4)$$

and

$$\frac{dF}{dE} = f_0 \cdot 10^{-11} \text{ TeV}^{-1} \text{ cm}^{-2} \text{ s}^{-1} \left(\frac{E}{E_0} \right)^{-\alpha} \exp\left(\frac{-E}{\beta}\right), \quad (5)$$

⁴ Table 6 is available only in the electronic edition of the journal, www.aanda.org

⁵ the respective log-P values are 83%, 48%, 73%, and 96%.

Table 3. Special fit results.

Data Set	Used Fit	E_0 [TeV]	f_0	α	β [TeV]	$\chi^2_{\text{red,fit}}$	Likelihood	E_{peak} [TeV]
2006/04/27	PL	0.5	9.54 ± 0.52	1.92 ± 0.07		5.3/4		
	log-P	0.5	11.5 ± 0.9	0.26 ± 0.17	1.2 ± 0.2	0.48/3	96%	
	PL+C	0.5	11.3 ± 1.2	1.44 ± 0.24	2.6 ± 1.3	0.34/3	96%	1.4 ± 1.0
All April Data	PL	0.5	4.53 ± 0.07	2.07 ± 0.04		16/5		
	log-P	0.5	4.84 ± 0.16	0.41 ± 0.11	0.69 ± 0.06	1.2/4	99%	
	PL+C	0.5	5.36 ± 0.31	1.77 ± 0.09	3.6 ± 1.1	1.8/4	99%	0.80 ± 0.42
High-State Nights	PL	0.5	8.19 ± 0.28	1.93 ± 0.05		6.0/4		
	log-P	0.5	9.21 ± 0.47	0.52 ± 0.17	1.1 ± 0.3	2.0/3	94%	
	PL+C	0.5	9.02 ± 0.64	1.75 ± 0.12	6.1 ± 4.0	3.1/3	91%	1.5 ± 1.2
Low-State Nights	PL	0.5	3.39 ± 0.10	2.17 ± 0.05		6.6/4		
	log-P	0.5	3.55 ± 0.13	0.41 ± 0.16	0.48 ± 0.12	1.1/3	97%	
	PL+C	0.5	4.15 ± 0.40	1.85 ± 0.15	2.9 ± 1.3	0.75/3	97%	0.45 ± 0.47

Results of a log-parabolic power-law fit in apex form (Eq. 6) and a power-law fit with an exponential cutoff ((Eq. 5) $\cdot (E/E_0)^2$) in $E^2 dF/dE$ after EBL de-absorption for special data sets. f_0 is given in units of $10^{-11} \text{TeV}^{-1} \text{cm}^{-2} \text{s}^{-1}$, α and β are the fit parameters as stated in the text, and Likelihood denotes the probability of a likelihood ratio test.

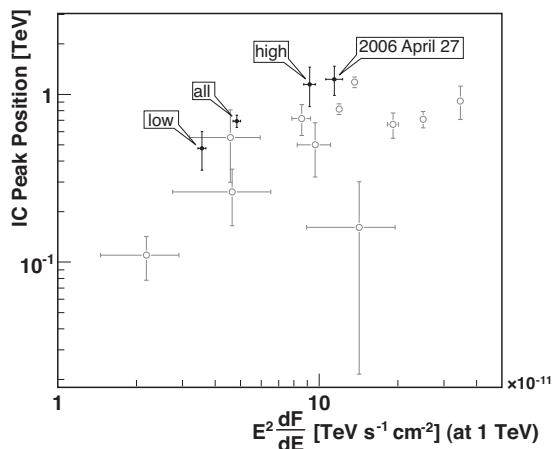


Fig. 5. Derived peak position using the log-P (Eq. 4) versus flux at 1 TeV for the data sets presented in Tab. 3. Historical data, taken from Albert et al. (2007a), are shown in gray. Our data confirm the indication of a correlation between the two parameters.

respectively. The likelihood ratio test results in a clear preference towards a log-P or a PL+C compared to a simple power-law with a probability of $\approx 96\%$ for both of them. The χ^2_{red} values for PL, log-P, and PL+C fits on the individual night-by-night spectra in Fig. 7 are given in Tab. 5.⁶ Also the high statistics data sets defined by combining all data from April, all data from the five low-state nights and all data from the three high-state nights, clearly showed evidence for a parabolic or cutoff shape of the spectra. The results of the fits and the probability of a likelihood ratio test are given in Tab. 3. For all these nights our data did not allow to prefer one model over the other. The fact that all of the high statistics data sets show a curved spectral shape is an indication of this feature being always visible for Mkn 421 and hence source intrinsic.

The curved power laws enable to locate a peak in the de-absorbed spectrum at $E_{\text{peak}} = E_0 10^{(2-\alpha)/(2\beta)}$ for the log-P and at $E_{\text{peak}} = (2-\alpha)\beta$ if $\alpha < 2$ for the PL+C fit. For simplicity

we determined E_{peak} of the log-P by using the apex form of the parabola in a logarithmic representation:

$$\log_{10} \frac{dF}{dE} = \log_{10}(f_0) + \log_{10}(\alpha) \left(\log_{10} \left(\frac{E}{0.5 \text{TeV}} / \frac{\beta}{0.5 \text{TeV}} \right) \right)^2 \quad (6)$$

which naturally yields both E_{peak} and the flux at the peak as β and f_0 , respectively. Additionally, the spectral cutoff is naturally obtained from the PL+C fit as the fit parameter β . The results are shown in Tab. 3. The values of E_{peak} as determined using the log-P and the PL+C were compatible with each other for the data sets averaging several nights and showed indications for an increase of the peak energy with rising flux level, as predicted if the VHE radiation were due to SSC mechanisms. We compare our results with historical values taken from Albert et al. (2007a) in Fig. 5. Our data confirm the previously suggested correlation.

The observation of a relation between flux (and thus, fluence) and the position of the VHE peak in the SED could be signalling a relation similar to the one suggested by Amati et al. (2002) and observed by Sakamoto et al. (2006) for gamma-ray bursts. Since the TeV γ -ray production is assumed to take place in a relativistic jet, and many of the same radiative processes are involved (on a larger scale, of course) it might be a similar (or related) mechanism at work on a different scale. A trend towards a relation between flux and spectral index in the TeV energy range has also been noted by Wagner (2008a), studying 17 known TeV blazars.

Although the peak energy measured on 2006 April 27 exceeds that of the All April Data and Low-State data set, it is, despite having a higher flux, comparable with that derived for the High-State data set. This discrepancy in terms of the expected behaviour in SSC models can be explained with the different nature of the data sets: The 2006 April 27 data represent a rather particular, 1.4 h long episode of an individual flare event, whereas the High-State data set is an average of three individual flares. Due to the sparse sampling, most probably each of these observations caught different epochs of the individual flare evolutions, during which the spectral shape can change considerably in terms of spectral index and curvature (see, e.g., Katarzyński et al., 2006). Hence the two data sets are not necessarily directly comparable.

The values of the derived cutoff energies are also suggesting this behavior, showing, with the exception of 2006 April 27, an increase with rising flux and indicating a source-intrinsic rather than a cosmological reason for the cutoff feature. This is in

⁶ Table 5 is available only in the electronic edition of the journal, www.aanda.org

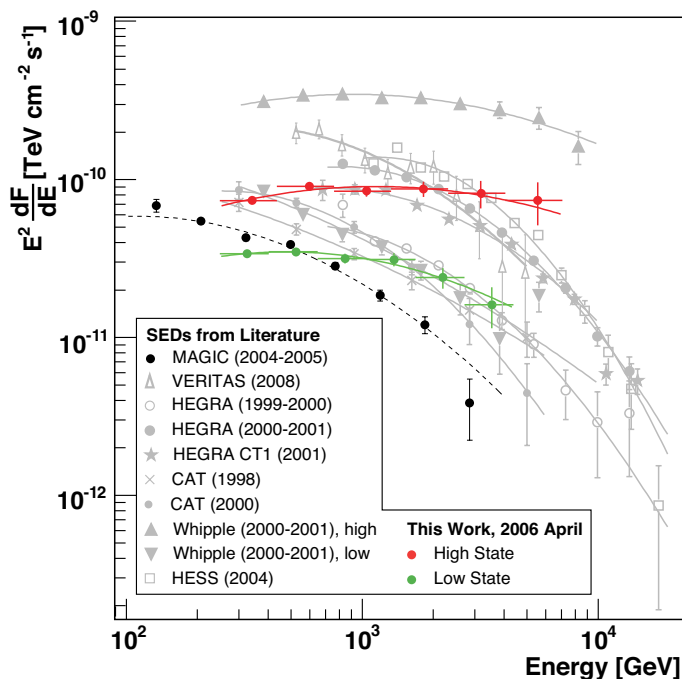


Fig. 6. EBL de-absorbed historical spectra of Mkn 421 (see Albert et al., 2007a, for references) along with selected spectra from the 2006 April campaign and the flare spectrum of Donnarumma et al. (2009). The solid line is the result of a fit using Eq. 4. Note that the historical data were deabsorbed using the model of Primack et al. (2005), our data and those from Donnarumma et al. (2009) with the model of Kneiske & Dole (2008).

accordance with the Kneiske & Dole (2008) lower-limit model, predicting an EBL cutoff for Mkn 421 at around 13 TeV.

In Fig. 6, we compare “historical” spectra measured between 1998 and 2005 with the low-state and high-state spectrum derived from the observations reported here. It is obvious that our low-state spectrum represents one of the lowest flux states ever measured in VHE for Mkn 421, whereas the high state spectrum shows no exceptionally high flux level of this source. Both spectra are harder than historical spectra with comparable flux levels, in particular harder than the VERITAS spectrum (Donnarumma et al., 2009), enabling one of the best measurements of the turnover of the SED in a low flux state. While a previous observation yielded a rather flat spectrum in the VHE regime (Aharonian et al., 2002), we conclude that we measured a rather clear peak (flat structure in the SED). The low-state spectrum has a shape similar to the one measured by HEGRA CT1, although at an approximately three times lower flux level. The high-state spectral shape resembles the high-state Whipple spectrum, which in turn has an about three times higher flux. This tendency can also be seen in Fig. 5, which shows that the fluxes we derive are systematically lower than historical measurements for comparable peak energies. Within the SSC framework this difference in flux for comparable spectral shapes can be caused by, e.g., a lower number of electrons with the same energy distribution as in the high-flux case.

In summary, we followed the evolution of a sequence of mild flares of the blazar Mkn 421 during one week from 2006 April 22 to 30, peaking at $F(E > 250\text{GeV}) = (3.21 \pm 0.15) 10^{-10} \text{ cm}^{-2}\text{s}^{-1}$ (≈ 2.0 Crab units). The nocturnal observations lasted at least for about one hour and allowed for the reconstruction of night-by-

night spectra. During three observation nights high fluxes were recorded, in which, however, no variability could be measured. In two of these nights, rather hard spectral indices were found, but this was also the case for the night with the lowest flux. During the night of 2006 April 29 with a not particularly high flux of $F(E > 250\text{GeV}) = (1.04 \pm 0.06) 10^{-10} \text{ cm}^{-2}\text{s}^{-1}$ (≈ 0.65 Crab units), clear intra-night variability with a flux-doubling time of $36 \pm 10_{\text{stat}}$ minutes was observed.

According to a likelihood ratio test, the spectra of some data sets were better described by curved power laws than simple power laws, enabling us to calculate peak and cutoff energies in the VHE regime. The derived peak values are consistent with an evolution of the peak energy with the flux, as suggested by historical data. Indications of an intrinsic cutoff in the spectra of Mkn 421, as found in former observations, are confirmed by our results.

During the *INTEGRAL*-triggered MWL campaign in 2006 June we observed Mkn 421 in one night at high zenith angles. Our measurements complement the three-night observations conducted by the Whipple 10-m telescope four days later. Taking the MAGIC and Whipple results together, a variability of Mkn 421 also during the *INTEGRAL* observations is evident. The energy coverage of the Whipple telescope spectrum ($\Delta E \approx 600$ GeV) was not sufficient to assess any spectral evolution by comparing it to the MAGIC spectrum ($\Delta E \approx 2$ TeV).

The determined fluxes and spectra will be further used for studies of the SED taking into account data taken at other photon energies in detailed MWL analyses (publications in preparation).

Acknowledgements. We thank the Instituto de Astrofísica de Canarias for the excellent working conditions at the Observatorio del Roque de los Muchachos in La Palma. The support of the German BMBF and MPG, the Italian INFN and Spanish MICINN is gratefully acknowledged. This work was also supported by ETH Research Grant TH 34/043, by the Polish MNiSzW Grant N N203 390834, and by the YIP of the Helmholtz Gemeinschaft.

References

- Acciari, V. A. et al. (VERITAS and MAGIC Collabs.), 2009, *ApJ*, 703, 169
 Aharonian, F. et al. (HEGRA Collab.), 2002, *A&A*, 393, 89
 Aharonian, F. et al. (HEGRA Collab.), 2004, *ApJ*, 614, 897
 Aharonian, F. et al. (H.E.S.S. Collab.), 2005, *A&A*, 437, 95
 Albert, J., et al. (MAGIC Collab.), 2006, *ApJ*, 638, L101
 Albert, J., et al. (MAGIC Collab.), 2007a, *ApJ*, 663, 125
 Albert, J., et al. (MAGIC Collab.), 2007b, *Nucl. Instrum. Meth.*, A583, 494
 Albert, J., et al. (MAGIC Collab.), 2008a, *ApJ*, 674, 1037
 Albert, J., et al. (MAGIC Collab.), 2008b, *Nucl. Instrum. Meth.*, A588, 424
 Albert, J., et al. (MAGIC Collab.), 2008c, *Nucl. Instrum. Meth.*, A594, 407
 Aliu, E., et al. (MAGIC Collab.), 2009, *Astropart. Phys.*, 30, 293
 Amati, L., et al., 2002, *Astron. Astrophys.*, 390, 81
 Baixeras, C., et al. (MAGIC Collab.), 2004, *Nucl. Instrum. Meth.*, A518, 188
 Błażejowski, H., et al., 2005, *ApJ*, 630, 130
 Bretz, T. & Wagner, R., 2003, in *Proceedings of the 28th International Cosmic Ray Conference*, Tsukuba, Japan, 5, 2947
 Bretz, T. & Dörner, D. (MAGIC Collab.), 2008, *AIP Conf. Proc.* 1085, 664
 Cortina, J., et al. (MAGIC Collab.), 2005, in *Proceedings of the 29th International Cosmic Ray Conference*, Pune, India, 5, 359
 Cortina, J., Goebel, F. & Schweizer, T. (MAGIC Collab.), 2009, in *Proceedings of the 31st International Cosmic Ray Conference*, Łódź, Poland (preprint: arXiv:0907.1211 [astro-ph])
 Costamante, L. & Ghisellini, G., 2002, *A&A*, 384, 56
 Dam, A. et al. (HEGRA Collab.), 1997, *Astropart. Phys.*, 8, 1
 Donnarumma, I., et al. (AGILE Collab., GASP-WEBT Collab., MAGIC Collab., VERITAS Collab.), 2009, *ApJ*, 691, L13
 Fegan, D. J., 1997, *J. Phys. G*, 23, 1013
 Fegan, D. J. (VERITAS Collab.), 2008, in *Proceedings of the 30th International Cosmic Ray Conference*, Merida, Mexico, 3, 901
 Fomin, V. P. et al., 1994, *Astropart. Phys.*, 2, 137
 Fossati, G. et al., 2008, *ApJ*, 677, 906
 Gaidos, J. A. et al., 1996, *Nature*, 383, 319

- Goebel, F., et al. (MAGIC Collab.), 2008a, in Proceedings of the 30th International Cosmic Ray Conference, Merida, Mexico, 3, 1025
- Goebel F., et al. (MAGIC Collab.), 2008b, in Proceedings of the 30th International Cosmic Ray Conference, Merida, Mexico, 3, 1481
- Gould, R. J. & Schröder, G. P., 1966, *Phys. Rev. Lett.*, 16, 252
- Hauser, M. G. & Dwek, E., 2001, *ARA&A*, 39, 249
- Hillas, A. M., 1985, in Proceedings of the 19th International Cosmic Ray Conference, La Jolla, 3, 445
- Horan, D., et al. (Whipple Collab.), 2009, *ApJ*, 695, 596
- Hsu, C.-C., et al. (MAGIC Collab.), 2009, in Proceedings of the 31st International Cosmic Ray Conference, Łódź, Poland, (preprint: arXiv:0907.0893 [astro-ph])
- Katarzyński, K., et al., 2006, *A&A*, 453, 47
- Knapp, J. & Heck, D., 2004, *EAS Simulation with CORSIKA: A Users Manual*
- Kneiske, T. M. & Dole, H., 2008, *AIP Conf. Proc.* 1085, 620
- Krennrich, F. et al., 2002, *ApJ*, 575, L9
- Krawczynski, H. et al., 2001, *ApJ*, 559, 187
- Jansen, B., et al., 2001, *A&A*, 365, 1
- Lessard, R. W. et al., 2001, *Astropart. Phys.*, 15, 1
- Li, T.-P. & Ma, Y.-Q., 1983, *ApJ*, 272, 317
- Lichti, G. G. et al., 2008, *A&A*, 486, 721
- Mannheim, K. et al., 1996, *A&A*, 315, 77
- Majumdar, P. et al., 2005, in Proceedings of the 29th International Cosmic Ray Conference, Pune, India, 5, 203
- Marscher, A. P., Gear, W. K., 1985, *ApJ*, 298, 11
- Maraschi, L., Ghisellini, G., Celotti, A., 1992, *ApJ*, 397, L5
- Massaro, E., et al., 2004, *A&A*, 413, 489
- Mazin, D., Goebel, F., Horns, D., Rowell, G., Wagner, R. M., Wagner, S., 2005, in Proceedings of the 29th International Cosmic Ray Conference, Pune, India, 4, 331
- Mazin, D. & Goebel, F., 2007, *ApJ*, 655, L13
- Meucci, M., et al., 2007, *IEEE Trans. Nucl. Sci.*, 54, 404
- Mitsuda, K., et al., 2007, *PASJ*, 59, S1
- Mücke, A. et al., 2003, *Astropart. Phys.*, 18, 593
- Nikishov, A. I., 1962, *Sov. Phys. JETP*, 14, 393
- Petry, D. et al., 1996, *A&A*, 311, L13
- Primack, J., et al., 2005, *AIP Conf. Proc.* 745, 23
- Punch, M. et al., 1992, *Nature*, 358, 477
- Riegel, B., et al. (MAGIC Collab.), 2005, in Proceedings of the 29th International Cosmic Ray Conference, Pune, India, 5, 215
- Sakamoto, T. et al., 2006, *Astrophys. J.*, 636, L73
- Schweizer, T., Wagner, R. M., Lorenz, E., 2008, *AIP Conf. Proc.* 1085, 455
- Steele, D. et al. (VERITAS Collab.), 2008, in Proceedings of the 30th International Cosmic Ray Conference, Merida, Mexico, 3, 989
- Tluczykont M., Shayduk M., Kalekin O., Bernardini E., 2007, *J. Phys. Conf. Ser.*, 60, 318
- Tonello, N., 2006, Ph.D. thesis, Technische Universität München, MPP-2006-21
- Wagner, R. M., 2008a, *MNRAS*, 385, 119
- Wagner, R. M., 2008b, *PoS(BLAZARS2008)*, 63, 13 (arXiv:0808.2483 [astro-ph])
- Yadav, K. K. et al. (TACTIC Collab.), 2007, *Astropart. Phys.*, 27, 447

¹ IFAE, Edifici Cn., Campus UAB, E-08193 Bellaterra, Spain

² ETH Zurich, CH-8093 Switzerland

³ INAF National Institute for Astrophysics, I-00136 Rome, Italy

⁴ Universidad Complutense, E-28040 Madrid, Spain

⁵ Technische Universität Dortmund, D-44221 Dortmund, Germany

⁶ Universitat Autònoma de Barcelona, E-08193 Bellaterra, Spain

⁷ Università di Padova and INFN, I-35131 Padova, Italy

⁸ Inst. de Astrofísica de Canarias, E-38200 La Laguna, Tenerife, Spain

⁹ University of Łódź, PL-90236 Łódź, Poland

¹⁰ Tuorla Observatory, University of Turku, FI-21500 Piikkiö, Finland

¹¹ Deutsches Elektronen-Synchrotron (DESY), D-15738 Zeuthen, Germany

¹² Max-Planck-Institut für Physik, D-80805 München, Germany

¹³ Università di Siena, and INFN Pisa, I-53100 Siena, Italy

¹⁴ Universitat de Barcelona (ICC/IEEC), E-08028 Barcelona, Spain

¹⁵ Universität Würzburg, D-97074 Würzburg, Germany

¹⁶ Depto. de Astrofísica, Universidad, E-38206 La Laguna, Tenerife, Spain

¹⁷ Università di Udine, and INFN Trieste, I-33100 Udine, Italy

¹⁸ Institut de Ciències de l'Espai (IEEC-CSIC), E-08193 Bellaterra, Spain

¹⁹ Inst. de Astrofísica de Andalucía (CSIC), E-18080 Granada, Spain

²⁰ Croatian MAGIC Consortium, Institute R. Boskovic, University of Rijeka and University of Split, HR-10000 Zagreb, Croatia

²¹ University of California, Davis, CA-95616-8677, USA

²² Inst. for Nucl. Research and Nucl. Energy, BG-1784 Sofia, Bulgaria

²³ INAF/Osservatorio Astronomico and INFN, I-34143 Trieste, Italy

²⁴ ICREA, E-08010 Barcelona, Spain

²⁵ Università di Pisa, and INFN Pisa, I-56126 Pisa, Italy

²⁶ supported by INFN Padova

²⁷ now at: Centro de Investigaciones Energéticas, Medioambientales y Tecnológicas (CIEMAT), Madrid, Spain

²⁸ now at: Max-Planck-Institut für Kernphysik, D-69029 Heidelberg, Germany

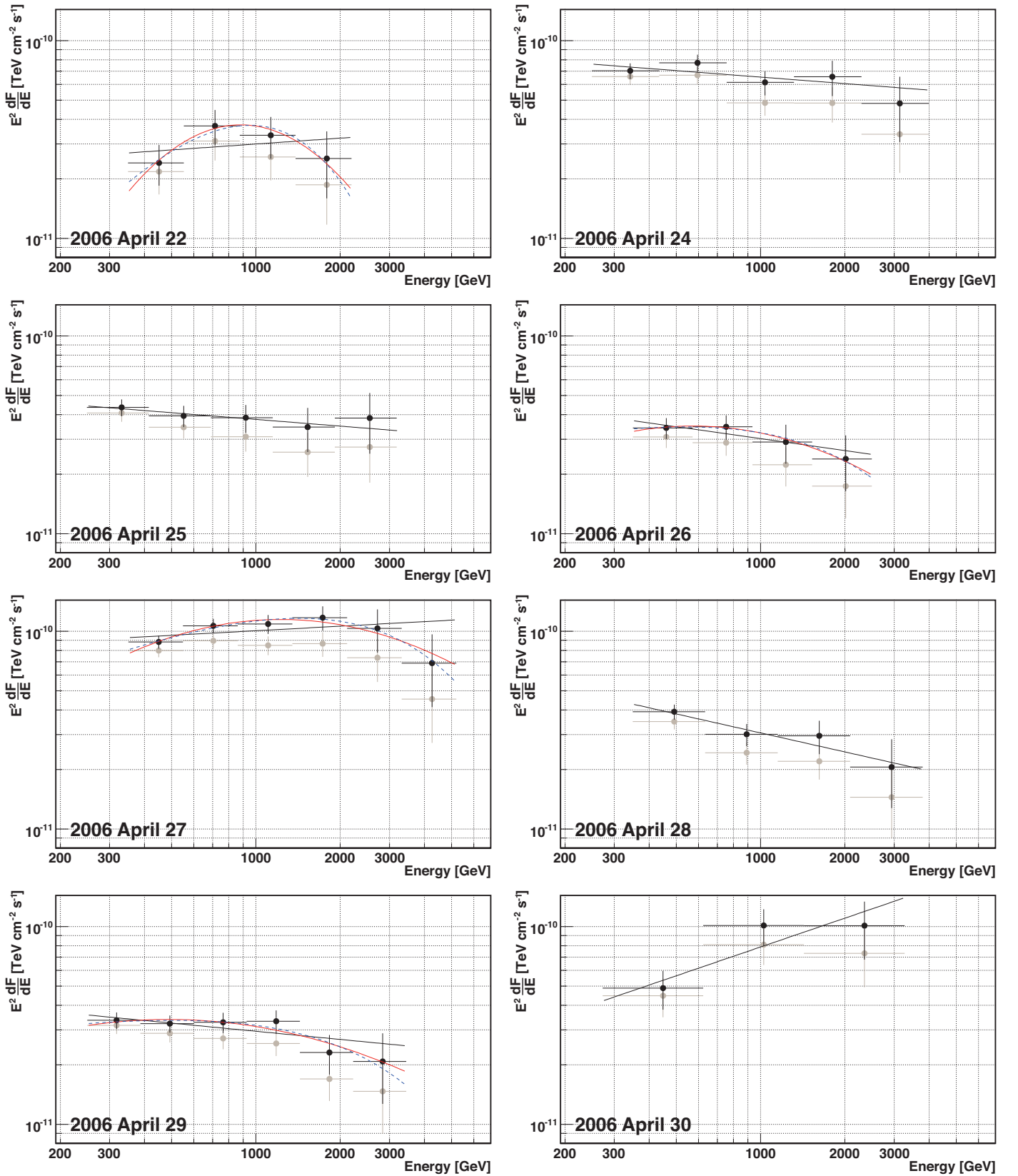


Fig. 7. Differential energy spectra for Mkn 421 for April 2006 before (gray points) and after (black points) correcting for EBL absorption. For the apparently hard spectra on 2006 April 22, 26, 27, and 29, log-P (Eq. 4) and PL+C (Eq. 5) fits were performed (red solid and blue dashed curves, respectively).

Table 4. Light curve data.

Observation [MJD]	$F_{>250\text{ GeV}}$ [$10^{-10}\text{ cm}^{-2}\text{ s}^{-1}$]
2006/04/22	
53847.98307	1.00 ± 0.21
53847.99775	0.70 ± 0.23
53848.00867	0.99 ± 0.17
2006/04/24	
53849.97136	2.56 ± 0.25
53849.98618	2.04 ± 0.21
53850.00033	2.37 ± 0.22
2006/04/25	
53850.93996	1.24 ± 0.15
53850.96431	1.49 ± 0.15
53850.98652	1.26 ± 0.15
2006/04/26	
53851.93677	0.97 ± 0.20
53851.95255	1.04 ± 0.21
53851.96726	1.25 ± 0.20
53851.98190	1.00 ± 0.18
53851.99680	1.13 ± 0.18
2006/04/27	
53852.94098	3.01 ± 0.33
53852.95502	3.19 ± 0.38
53852.96823	3.05 ± 0.31
53852.98159	3.57 ± 0.32
53852.99406	3.17 ± 0.28
2006/04/28	
53853.88754	1.36 ± 0.25
53853.89880	0.80 ± 0.27
53853.92893	1.22 ± 0.24
53853.93984	1.09 ± 0.25
53853.95457	1.18 ± 0.24
53853.96887	1.22 ± 0.27
53853.98040	0.95 ± 0.19
53853.99316	1.32 ± 0.21
53854.00687	1.12 ± 0.18
2006/04/29	
53854.90199	2.07 ± 0.23
53854.91620	1.42 ± 0.19
53854.95206	0.86 ± 0.16
53854.96625	1.11 ± 0.17
53854.97974	1.27 ± 0.18
53854.99354	0.83 ± 0.15
53855.00847	0.80 ± 0.14
53855.02879	0.69 ± 0.11
2006/04/30	
53855.97595	2.39 ± 0.33
2006/06/14	
53900.92797	0.45 ± 0.09
53900.94585	0.26 ± 0.08

Table 6. Energy spectra for all observation nights under study after EBL de-absorption.

E bounds [GeV]		Flux [$\text{TeV cm}^{-2}\text{ s}^{-1}$]
2006/04/22		
350	554	$(2.39 \pm 0.56) \cdot 10^{-7}$
554	877	$(3.67 \pm 0.75) \cdot 10^{-7}$
877	1389	$(3.29 \pm 0.79) \cdot 10^{-7}$
1389	2200	$(2.51 \pm 0.93) \cdot 10^{-7}$
2006/04/24		
250	435	$(7.00 \pm 0.64) \cdot 10^{-7}$
435	758	$(7.69 \pm 0.76) \cdot 10^{-7}$
758	1320	$(6.12 \pm 0.86) \cdot 10^{-7}$
1320	2297	$(6.54 \pm 1.33) \cdot 10^{-7}$
2297	4000	$(4.80 \pm 1.74) \cdot 10^{-7}$
2006/04/25		
250	416	$(4.36 \pm 0.43) \cdot 10^{-7}$
416	693	$(3.95 \pm 0.48) \cdot 10^{-7}$
693	1154	$(3.86 \pm 0.62) \cdot 10^{-7}$
1154	1922	$(3.47 \pm 0.87) \cdot 10^{-7}$
1922	3200	$(3.85 \pm 1.31) \cdot 10^{-7}$
2006/04/26		
350	572	$(3.41 \pm 0.41) \cdot 10^{-7}$
572	935	$(3.46 \pm 0.49) \cdot 10^{-7}$
935	1529	$(2.90 \pm 0.65) \cdot 10^{-7}$
1529	2500	$(2.38 \pm 0.75) \cdot 10^{-7}$
2006/04/27		
350	549	$(8.83 \pm 0.66) \cdot 10^{-7}$
549	860	$(1.07 \pm 0.09) \cdot 10^{-6}$
860	1349	$(1.09 \pm 0.12) \cdot 10^{-6}$
1349	2115	$(1.17 \pm 0.17) \cdot 10^{-6}$
2115	3317	$(1.03 \pm 0.26) \cdot 10^{-6}$
3317	5200	$(6.89 \pm 2.76) \cdot 10^{-7}$
2006/04/28		
350	635	$(3.92 \pm 0.34) \cdot 10^{-7}$
635	1153	$(3.01 \pm 0.39) \cdot 10^{-7}$
1153	2093	$(2.96 \pm 0.57) \cdot 10^{-7}$
2093	3800	$(2.06 \pm 0.79) \cdot 10^{-7}$
2006/04/29		
250	387	$(3.37 \pm 0.32) \cdot 10^{-7}$
387	600	$(3.24 \pm 0.32) \cdot 10^{-7}$
600	929	$(3.29 \pm 0.39) \cdot 10^{-7}$
929	1438	$(3.33 \pm 0.45) \cdot 10^{-7}$
1438	2228	$(2.31 \pm 0.53) \cdot 10^{-7}$
2228	3450	$(2.08 \pm 0.81) \cdot 10^{-7}$
2006/04/30		
250	572	$(4.87 \pm 1.09) \cdot 10^{-7}$
572	1310	$(1.01 \pm 0.21) \cdot 10^{-6}$
1310	3000	$(1.01 \pm 0.33) \cdot 10^{-6}$
2006/06/14		
450	669	$(2.76 \pm 0.74) \cdot 10^{-7}$
669	995	$(1.61 \pm 0.67) \cdot 10^{-7}$
995	1480	$(2.54 \pm 0.79) \cdot 10^{-7}$
1480	2200	$(1.80 \pm 0.85) \cdot 10^{-7}$

The two energy bounds specify the range in which the corresponding flux was measured.

Table 5. χ^2_{red} values for the PL, log-P, and PL+C fits performed in Fig. 7

	22	24	25	26	27	28	29	30
PL	2.2/2	1.9/3	0.21/3	0.47/2	5.3/4	0.59/2	2.3/4	1.5/1
log-P	0.14/1			0.041/1	0.48/3		1.1/3	
PL+C	0.27/1			0.076/1	0.34/3		0.87/3	

The columns represent days in 2006 April.